Constraints on Nucleon Decay via “Invisible” Modes from the Sudbury Neutrino Observatory

(SNO Collaboration)

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Data from the Sudbury Neutrino Observatory have been used to constrain the lifetime for nucleon decay to “invisible” modes, such as \( n \rightarrow 3\nu \). The analysis was based on a search for \( \gamma \)-rays from the de-excitation of the residual nucleus that would result from the disappearance of either a proton or neutron from \( ^{16}\text{O} \). A limit of \( \tau_{\text{inv}} > 2 \times 10^{29} \) years is obtained at 90% confidence for either neutron or proton decay modes. This is about an order of magnitude more stringent than previous constraints on invisible proton decay modes and 400 times more stringent than similar neutron modes.


Experimental signatures of the grand unification of the electroweak and strong interactions have been sought with increasing sensitivity for the past twenty-five years. Much effort has gone to identifying specific decay modes of free protons and bound nucleons as signatures of grand unification, but no signal has been observed to date. Decay modes more unusual than those typically explored cannot, however, be ruled out (see, for example, \[1\]). A recent paper has even suggested a model in which \( n \rightarrow 3\nu \) becomes the dominant mode \[2\]. Thus, the search for any mode which may have been missed by previous experiments is of fundamental interest. In this paper, the Sudbury Neutrino Observatory (SNO) \[3\] is used to search for what we will refer to as “invisible” decay channels,
i.e., those in which no visible energy is deposited in the detector via direct production of energetic, charged particles.

The search utilizes SNO’s unique detection capabilities for low energy γ-rays, based on the Cherenkov light produced by the resulting Compton-scattered electron. A generic signature for nucleon disappearance in \(^{16}\)O arises from the subsequent de-excitation of the residual nucleus [4,5]. Approximately 45% of the time, the de-excitation of either \(^{15}\)O* or \(^{15}\)N\(^+\) results in the production of a γ-ray of energy 6-7 MeV. SNO detects these γ-rays with good efficiency. In fact, the primary energy calibration source used by SNO is the 6.13-MeV γ-ray produced in the de-excitation of \(^{16}\)N.

A background to this nucleon-decay signal in SNO results from neutral current (NC) interactions of solar neutrinos. This is due to the γ-rays of similar energies that are produced as a result of neutron captures on nuclei in the detector. In Phase I (D\(_2\)O), neutrons were detected through observation of the single 6.25-MeV γ-ray resulting from neutron capture on deuterium. This capture efficiency is 0.29, but the threshold applied to the analysis to limit low energy backgrounds reduces the overall neutron detection efficiency to \(\epsilon_n = 0.144 \pm 0.005\) [6]. In Phase II (D\(_2\)O+NaCl), two tonnes of NaCl were added to the one kilotonne of D\(_2\)O. Neutron captures on \(^{35}\)Cl release 8.6 MeV of energy in γ-rays, with most capture events producing multiple γ-rays. The corresponding capture efficiency in Phase II is 0.90 and, due to a relatively high analysis threshold, the overall neutron detection efficiency is \(\epsilon'_n = 0.399 \pm 0.016\) [6]. The multiple γ-rays from neutron captures in Phase II result in a more isotropic distribution of Cherenkov light, which can be used as a further discriminant for identifying the neutron-induced component. However, the principal advantage in comparing Phase I and Phase II data lies in the fact that γ-rays from the nucleon-decay signal are detected with similar efficiencies in SNO, while neutrons produced by \(^{8}\)B solar neutrinos are detected with very different efficiencies. These characteristics are used in what follows to measure an upper limit for nucleon disappearance.

In terms of the data from Phase I of SNO, the rate of nuclear γ-ray production can be related to the apparent production rate of neutrons by taking account of the detection efficiencies for neutrons, γ-rays and particle misidentification as follows:

\[
R_{\gamma}\epsilon_n f_{\gamma n} = R_n - \epsilon_n R_{NC}
\]

where \(R_{\gamma}\) is the rate of nuclear γ-ray production due to nucleon decay; \(\epsilon_n\) is the efficiency for detecting the nuclear γ-rays above the analysis energy threshold; \(f_{\gamma n}\) is the fraction of the detected nuclear γ-rays which are mistaken for neutrons; \(R_n\) is the extracted neutron detection rate nominally attributed to NC interactions; \(\epsilon_n\) is the neutron detection efficiency for the fiducial volume and analysis energy threshold; and \(R_{NC}\) is the actual production rate of neutrons due to solar neutrino NC interactions. Similarly, for Phase II data,

\[
R_{\gamma}\epsilon'_n f'_{\gamma n} = R'_n - \epsilon'_n R_{NC}
\]

Thus,

\[
R_{\gamma} = \frac{R_n - \epsilon_n R_{NC}}{\epsilon_n f_{\gamma n} - \epsilon'_n f'_{\gamma n}} \equiv \frac{\Delta R_n}{\epsilon_n f_{\gamma n} - \epsilon'_n f'_{\gamma n}}
\]

where \(\Delta R_n\) is the difference between the extracted neutron detection rate attributed to NC interactions in Phase I and that implied by data from Phase II.

In order to compare Phase I and Phase II rates under the same assumption for the underlying CC spectrum, results from SNO data were used in which the CC component was constrained to follow the shape of a standard \(^{8}\)B energy spectrum [7]. Table I summarizes the relevant results from these two phases. The extracted numbers of CC, NC and ES (elastic scattering) events include the subtraction of all known backgrounds (as detailed in [6] and [8]), including atmospheric neutrino interactions which might identically mimic the nucleon-decay signal via the removal of a proton or neutron from \(^{16}\)O.

The number of NC events extracted in Phase I [6] was 576.5 for a livetime of 306.4 days, yielding a rate of \(R_n = 686.8 \pm 83.9\) per year, with statistical and systematic uncertainties added in quadrature. Similarly, for Phase II, 1265.8 NC events were implied based on a livetime of 254.2 days, yielding a rate of \(R'_n = 1817.6 \pm 136.6\) per year. Thus, accounting for the relative neutron detection efficiencies,

\[
\Delta R_n = \frac{R_n - \epsilon_n R_{NC}}{\epsilon'_n f'_{\gamma n} - \epsilon_n f_{\gamma n}}
\]

\[
= \frac{R_n - \epsilon_n R_{NC}}{\epsilon'_n f'_{\gamma n} - \epsilon_n f_{\gamma n}}
\]

\[
= \frac{(686.8 \pm 83.9) - (656.0 \pm 49.3)}{30.8 \pm 97.3}
\]

Thus, an upper limit of \(\Delta R_n < 180.6\) per year at 90% confidence limit is obtained using a standard, Bayesian prescription (which is also in good agreement with frequentist prescriptions) [9].

From [9], a vanishing neutron from the \(^{16}\)O nucleus results in an excited state which has a branching ratio of 44% for producing a 6.18-MeV γ and 2% for a 7.03-MeV γ. For a vanishing proton, the distribution is nearly the same, with a branching ratio of 41% for a 6.32-MeV γ and 4% for a 7.0-MeV γ. The signal extraction procedures previously used for solar neutrino analyses were applied to simulated nuclear γ-ray lines of these energies, combined with a simulated solar neutrino signal. The numbers of additional NC events extracted relative to the actual NC signals generated were then expressed as fractions of the generated nuclear γ-ray signals. The values of \(f_{\gamma n}\) and \(f'_{\gamma n}\) were then determined by the appropriate weighting of these fractions in accordance with the relative branching ratios given above. For Phase I
data, it was found that $f_{\gamma n} = 0.99^{+0.01}_{-0.02}$ for both neutron and proton decay modes. This is as expected since the neutron signal in pure D$_2$O results from a 6.25-MeV $\gamma$-ray, which is virtually indistinguishable from either 6.18 MeV or 6.32 MeV within the energy resolution of the detector. The distributions are, therefore, nearly 100% covariant. For Phase II data, $f_{\gamma n} = 0.75_{-0.01}^{+0.01}$ (again, nearly identical for either decay mode). Once more, this is roughly what is expected given the additional isotropy information. The lower value of $f_{\gamma n}$ derived reflects a compromise within the fitting procedure between providing a good description of the isotropy distribution and the energy spectrum expected for neutrons. These same simulated nuclear $\gamma$-ray lines were also used to determine $\epsilon_\gamma$ and $\epsilon_\gamma'$. For neutron (proton) decay modes, these were found to be $0.51 \pm 0.01$ ($0.59 \pm 0.01$) and $0.361 \pm 0.005$ ($0.425 \pm 0.006$), respectively.

Thus, an upper limit can be deduced for the number of decay $\gamma$-rays at greater than 90% confidence level of $R_{5\text{im}}^{\text{lim}} < 443$ per year for neutron decay and $R_{5\text{im}}^{\text{lim}} < 385$ per year for proton decay. An upper bound to invisible modes of nucleon decay can now be established as follows:

$$\tau_{\text{inv}} > \frac{N_{np}}{R_{5\text{im}}^{\text{lim}}} \epsilon_\gamma$$

where $N_{np}$ is the number of neutrons or protons (depending on decay mode) within the D$_2$O fiducial volume which are bound in $^{16}$O ($1.85 \times 10^{22}$), and $\epsilon_\gamma$ is the efficiency for the decay to result in the release of a 6 or 7-MeV $\gamma$-ray (0.46 for neutron modes and 0.45 for proton modes). Therefore, the comparison of Phase I and Phase II data from SNO implies that, at greater than 90% confidence level,

$$\text{for neutron modes: } \tau_{\text{inv}} > 1.9 \times 10^{29} \text{ years}.$$  

$$\text{for proton modes: } \tau_{\text{inv}} > 2.1 \times 10^{29} \text{ years}.$$  

Prior to this paper, the best constraint on $n \rightarrow 3\nu$ used by the Particle Data Group was based on Kamiokande data in which higher energy, but much weaker, branches of the de-excitation of the oxygen nucleus were considered and yielded a limit of $\tau > 5 \times 10^{26}$ years. It has been proposed that a similar analysis could be carried out with data from Super-Kamiokande and, by making use of the carbon nucleus, possibly even in the KamLAND detector. It has also been noted that the disappearance of a proton from the deuteron in heavy water detectors would result in a free neutron, which could then be captured to yield a detectable signal for invisible proton decay (see, for example [11]). This has already been used to yield a lower bound to the proton lifetime in excess of $10^{28}$ years for such modes [6,12]. Lead percolate has also been suggested as a possible future detector medium to search for invisible nucleon decay, making use of de-excitation of the nuclear hole that would be left in $^{35}$Cl, with an estimated sensitivity on the order of $10^{30}$ years for a one kilotonne detector [13,14]. Owing to the extremely low background levels in SNO, the principal branches of the de-excitations for $^{16}$O have been probed here and have yielded limits which are within a factor of 5 of this level. Thus, the constraint presented here is about an order of magnitude more stringent than the recently published limits on invisible proton-decay and 400 times more stringent than previous limits on neutron modes, such as $n \rightarrow 3\nu$.

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**TABLE I: Signal extraction results for CC constrained to $^8$B shape.** Error bars are the quadrature sum of statistical and systematic uncertainties.

<table>
<thead>
<tr>
<th>Analysis Parameter</th>
<th>Phase I (pure D$_2$O)</th>
<th>Phase II* (D$_2$O + NaCl)</th>
</tr>
</thead>
<tbody>
<tr>
<td>Fiducial Volume</td>
<td>$6.97 \times 10^9$ cm$^3$</td>
<td>$6.97 \times 10^9$ cm$^3$</td>
</tr>
<tr>
<td>Energy Threshold</td>
<td>$T_{\text{eff}} &gt; 5$ MeV</td>
<td>$T_{\text{eff}} &gt; 5.5$ MeV</td>
</tr>
<tr>
<td>Livetime</td>
<td>306.4 days</td>
<td>254.2 days</td>
</tr>
<tr>
<td>CC Events</td>
<td>1967.7±117.9</td>
<td>1430.3±97.1</td>
</tr>
<tr>
<td>ES Events</td>
<td>263.6±29.2</td>
<td>163.7±23.8</td>
</tr>
<tr>
<td>NC Events</td>
<td>576.5±70.4</td>
<td>1265.8±95.2</td>
</tr>
<tr>
<td>Neutron Detection Efficiency</td>
<td>0.144±0.005</td>
<td>0.399±0.010</td>
</tr>
<tr>
<td>NC Event Rate ($R_n$ &amp; $R_{n^*}$)</td>
<td>686.8±83.9 yr$^{-1}$</td>
<td>1817.6±136.6 yr$^{-1}$</td>
</tr>
<tr>
<td>Equivalent Phase I NC Rate ($R_n$ &amp; $\frac{\epsilon_n R_{n^*}}{\epsilon_\gamma}$)</td>
<td>686.8±83.9 yr$^{-1}$</td>
<td>656.0±49.3 yr$^{-1}$</td>
</tr>
</tbody>
</table>

* The Phase II data set used for this analysis is identical to that presented in [11]
4

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[8] By contrast, the main thrust of [7] was a measurement of the total solar neutrino flux independent of assumptions regarding the CC spectral shape.
[14] In principle, de-excitation of holes left in $^{35}$Cl can also be explored in SNO from Phase II data, in which NaCl was added. However, the concentration of this was only about 0.2% by weight, so the sensitivity to nucleon decay is poor relative to $^{16}$O.